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For this longitude,  $151^{\circ} 48'$  W., the corresponding latitude of the center of the eclipse path is  $11^{\circ} 17'$  S. The island thus appears to be within ten miles of the central line. As the shadow-path will be about ninety-three miles wide, this location would be comparatively close to the center. It is not probable that the assigned latitude is seriously in error. Any possible error in the assigned longitude would hardly be serious, for the shadow will travel nearly east and west. The motion in longitude will be six times that in latitude.

The duration of totality at Flint Island, from American Ephemeris data, would be  $4^m 0^s$ .

The next total eclipse of the Sun observable in the United States will take place on June 8, 1918. It will pass from the northwest coast of Washington to Florida during the latter half of the summer afternoon.

Another total eclipse, occurring on September 10, 1923, will enter the coast of California at latitude about  $36^{\circ}$  in the early afternoon, and will pass over the Gulf of Mexico and southwestern Cuba.

W. W. CAMPBELL.

#### NOTE ON COMET *c* 1905 (GIACOBINI).

The third comet of the year 1905 was discovered December 6th by GIACOBINI at Nice. At the time of discovery it was  $21^{\circ}$  north of the equator in Right Ascension  $14^h 22^m$ . Since then it has been traveling southeasterly toward the Sun. Its closest approach to the Sun will be on the 22d of this month, when it will be within twenty million miles of that body. Its closest approach to the Earth was on the 6th of January, when it was distant about one hundred million miles. The plane of its orbit is inclined  $43^{\circ}$  to the plane of the ecliptic. The orbit is parabolic.

Because of the very near approach to the Sun this comet becomes very bright at the end of January. At the present date it is visible to the unaided eye a few moments before sunrise. It will pass the Sun and become an evening object on the 23d inst. By the 28th it will set about half an hour after the Sun, and will be  $7^{\circ}$  south of it. On this date it should be easily seen with the unaided eye, when it will be thirty times brighter than at discovery. After the first of February its brightness will diminish rapidly.

Two sets of elements and ephemerides for this comet have been derived by LEUSCHNER's "short method" by the writer, with the assistance of Mr. A. J. CHAMPREUX. These may be found in Lick Observatory *Bulletins* Nos. 87 and 88.

RUSSELL TRACY CRAWFORD.

BERKELEY ASTRONOMICAL DEPARTMENT, January 17, 1906.

#### PHOTOGRAPHS OF COMET *c* 1905.

Comet *c* 1905 was photographed at the Lick Observatory on nine different dates preceding its perihelion passage. The first photograph was taken on December 23, 1905, when the comet was comparatively faint and no suggestion of a tail was visible in the observing telescope. When the plate was developed the comet was found to have a tail extending several degrees from the nucleus. On account of stormy weather it was impossible to secure another photograph until December 28th. By this time the comet had become very much brighter, the tail being easily visible in the guiding telescope. This plate showed a tail extending a distance of eight or ten degrees from the nucleus, and much detail in its structure was brought out.

Subsequent plates recorded many changes in the structure of the tail, and gave evidence of rapid motion in the material composing it. The exposure times varied from half an hour to an hour, depending upon weather conditions and the position of the comet.

Besides the plates described above a series of trail plates was taken, from which it is hoped data of value may be obtained concerning the variation in brightness of the comet's nucleus.

At a later date the plates will be studied in detail. In the mean time a more extensive series may be secured. The comet has passed perihelion, and will soon be an evening object for telescopic observation.

January, 1906.

ELLIOTT SMITH.

#### NOTE ON THE COMETS DISCOVERED AT THE LOWELL OBSERVATORY.

A telegram was received at the Lick Observatory on the evening of December 14, 1905, from Professor E. C. PICKERING